

Densities and temperatures in the Venus mesosphere and lower thermosphere retrieved from SOIR on board Venus Express: Retrieval technique

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[1] The SOIR instrument, flying on board Venus Express, operates in the infrared spectral domain and uses the solar occultation technique to determine the vertical profiles of several key constituents of the Venus atmosphere. The retrieval algorithm is based on the optimal estimation method, and solves the problem simultaneously on all spectra belonging to one occultation sequence. Vertical profiles of H₂O, CO, HCl, and HF, as well as some of their isotopologues, are routinely obtained for altitudes ranging typically from 70 to 120 km, depending on the species and the spectral region recorded. In the case of CO₂, a vertical profile from 70 up to 150 km can be obtained by combining different spectral intervals. Rotational temperature is also retrieved directly from the CO₂ signature in the spectra. The present paper describes the method used to derive the above mentioned atmospheric quantities and temperature profiles. The method is applied on some retrieval cases illustrating the capabilities of the technique. More examples of results will be presented and discussed in a following companion paper which will focus on the CO₂ vertical profiles of the whole data set.

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1. Introduction

[2] Venus is known to be a very warm planet with a dense atmosphere. Its mesosphere extending from the cloud deck up to \sim 100 km is composed mainly of carbon dioxide (CO₂, 96.5%) and Nitrogen (N₂, 3.5%). This region is dynamically and photochemically very active, but has been up to now poorly observed. Previous measurements indeed focused essentially on the regions below the clouds down to the surface. The mesosphere acts as a transition region between the lower atmosphere, and its cloud deck and superrotational zonal winds, and the thermosphere, dominated by the subsolar to antisolar circulation. The altitude of 100 km is a natural boundary in the atmosphere of Venus. The information above 100 km altitude about temperature, pressure, the densities of carbon dioxide and other minor atmospheric constituents, their sources and sinks, as well as the reactions in which they are involved, is incomplete, as only scarce measurements have been performed [Clancy *et al.*, 2003; Sandor and Clancy, 2005;

Clancy *et al.*, 2008]. Chemically active species, such as sulphur-bearing gases (COS, SO₂) and halides (HCl, HF) have already been reported [*de Bergh et al.*, 2006] but below 70 km altitude.

[3] The Solar Occultation in the Infrared (SOIR) spectrometer is an extension of the French instrument SPICAV, the combination being called SPICAV/SOIR [Bertaux *et al.*, 2007; Svedhem *et al.*, 2007]. SPICAV/SOIR is one of the seven instruments on board Venus Express (VEX), a planetary mission of the European Space Agency (ESA) that was launched in November 2005 and inserted into a polar orbit around Venus in April 2006 [Titov *et al.*, 2006]. VEX is orbiting around Venus on a polar orbit, which allows the instrument to sound almost all latitudes. The region comprised between 40° and 60° North is however more difficult to sound due to the orbit shape.

[4] SOIR is the first high-resolution near-infrared spectrometer on board a spacecraft investigating the Venustian mesosphere where it detected H₂O and HDO [Fedorova *et al.*, 2008], SO₂ [Belyaev *et al.*, 2008], HCl, HF, CO, and CO₂ [Vandaele *et al.*, 2008; Belyaev *et al.*, 2008; Fedorova *et al.*, 2008; Vandaele *et al.*, 2008]. The newness of such a spectrometer and the high accuracy obtained allow us to measure Venus' atmosphere minor constituents as well as CO₂ absorption lines with a never-so-far obtained resolution from orbit. The solar occultation technique is used for the first time at Venus. This technique has the advantage of being a

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Table 1. Summary of Pixel Rows Combinations

Denomination	Number of Scanned Orders per Second	Number of Lines in Each Group	Number of Lines Groups
2×16	4	16	2
2×12	4	12	2
4×4	2	4	4
4×3	2	3	4
8×4	1	4	8
8×3	1	3	8

self calibrated method, as will explained below in the text. The retrieval method developed in this study will be discussed in details in terms of sensitivity and error sources. The algorithm will be used to derive the atmospheric parameters from the spectra measured in two different regions during an occultation: CO₂, CO and temperature vertical profiles will be derived. We will also show that information on the temperature can be retrieved directly from the spectra. Comprehensive results considering CO₂ will be given in a companion paper (A. Mahieux et al., Densities and temperatures in the Venus mesosphere and lower thermosphere retrieved from SOIR on board Venus Express: CO₂ densities and temperatures, manuscript in preparation, 2010) in which the focus will be made on temporal as well as latitudinal distribution of the CO₂ density and temperature.

2. SOIR Instrument

[5] SOIR is an innovative, compact, high-resolution spectrometer for remote infrared measurements of gaseous species of the Venus' atmosphere, which has been already extensively described [Nevejans et al., 2006; Mahieux et al., 2008, 2009]. A short description of its working principles will be given here. The system consists of an echelle grating spectrometer working in the infrared (IR) combined with an acousto optic tunable filter (AOTF) for the selection of diffraction grating orders. Such an instrument combines the advantages of a high spectral resolution, fast and flexible access to any part of its spectral domain, compactness, no moving parts and low mass. The SOIR instrument is designed to measure at high resolution ($\sim 0.15 \text{ cm}^{-1}$) the atmospheric transmission in the IR, between 2.2 and 4.3 μm (2200–4400 cm^{-1}).

[6] The AOTF is driven by high radio frequencies (RF) that may be tuned for selecting the band-pass wave number range. The AOTF is a fast response solid state TeO₂ optical filter, whose transfer function has been extensively investigated by Mahieux et al. [2009] and may be approximated by the sum of sinc squared functions, where sinc(x) is defined as $\sin(\pi x)/\pi x$, x being the independent variable, here the RF. The main lobe of the sinc square-like function has a full width at half maximum (FWHM) of about 24 cm^{-1} . The shape of the AOTF transfer function varies significantly along the wave number range [Mahieux et al., 2009].

[7] The echelle grating of SOIR divides the spectral range into 94 useful diffraction orders, ranging from 101 to 194. The bandwidth varies from 19.3 cm^{-1} in diffraction order 101 to 37.1 cm^{-1} in order 194. The free spectral range (FSR) of SOIR equals 22.38 cm^{-1} . In diffraction orders 101 to 122, the FSR is greater than the bandwidth, and hence, the

detector will lose part of the spectrum. The opposite occurs in diffraction orders 123 to 194, where only a part of the detector will be covered by the diffracted order. Order overlapping occurs in these orders. The AOTF FWHM is greater than the FSR, which will induce an additional contribution of the adjacent orders to the scanned order. The resolution of SOIR varies from order to order, with a value of about 0.11 cm^{-1} in diffraction order 101 to 0.21 cm^{-1} in order 194.

[8] The slit of SOIR is rectangular, with a viewing angle of 2' in the spectral direction, and 30' in the spatial direction. The apparent size of the Sun at Venus is 44', which ensures that the slit will remain within the solar disc. The detector counts 320 pixels in the spectral direction and 256 pixels in the spatial direction called pixel rows further in the text. However, the light coming from the aperture only covers 32 pixel rows in the spatial direction. From onboard measurements, it has been observed that on these 32 pixel rows, the side rows have a very bad signal-to-noise ratio. For that reason only the 24 central pixel rows are read. This has been applied to measurements which occurred after orbit 332 (19 March 2007). The pixel field of view varies between 0.03 and 0.06 cm^{-1} as a function of the pixel position and the diffraction order.

[9] Due to telemetry constraints, SOIR can only send 8 measurement packets per second. The AOTF and the electronics of SOIR limit the number of orders to be measured to 4 orders per second. With all these constraints, six different cases have been defined to sound Venus' atmosphere. The pixel rows are summed up and groups of pixel rows have been defined. They are summarized in Table 1.

[10] As an example, and throughout the following text, we will consider the 2×12 pixel rows case. As 4 diffraction orders can be measured successively during one second, each order is scanned during 250 ms, and the same order is scanned once in every second. Two groups of lines are defined, each group results of the sum of 12 pixel rows on the detector. This operation improves also the signal-to-noise ratio. Typical values for the signal-to-noise ratio range between 500 and 3000. As the SOIR instrument is able to measure at best 4 diffraction orders almost simultaneously during one solar occultation, 4 different wave number ranges and therefore various species can be probed almost simultaneously.

3. Solar Occultation Technique

[11] The solar occultation technique used by SOIR allows the derivation of unique information about the vertical structure and composition of the Venus mesosphere [Vandaele et al., 2008]. SOIR is looking toward the Sun and records spectra on a one second cycle basis. Solar occultations occur when the line of sight of the instrument crosses the atmosphere of Venus. The projection of the center of the slit on the limbs during each single measurement defines the tangent altitude. Because the spacecraft is moving along its orbit, the instrument sounds the atmosphere of the planet at different tangent altitudes, see Figure 1. Two different configurations can be observed: the ingress case, when the tangent altitude decreases with time, and the egress case, when the tangent altitude increases with time. For the sake of simplicity, only the ingress case will be described here.

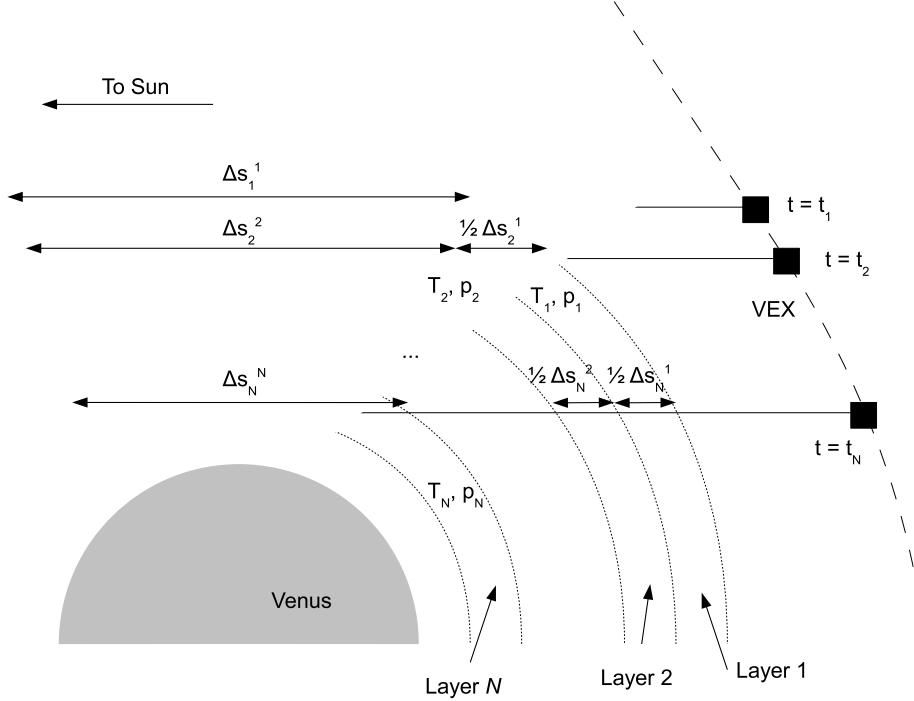


Figure 1. Onion-peeling definition while performing a solar occultation. The tangent altitude of each measurement defines a layer. The layers are spherical, and the atmospheric parameters such as the temperature T_i , pressure p_i , and partial pressure of a given species are considered to be constant within each layer i . The light path length within layer i of the measurement j (having its tangent altitude in layer j) is denoted Δs_i^j .

The egress case may be considered as ingress with the time scale reversed.

[12] The instrument is turned on and, after a precooling phase to ensure that the detector temperature has reached its operating value of about 70 K, it starts recording solar spectra before the line of sight crosses the atmosphere. It is assumed that the atmosphere of Venus extends up to 220 km in altitude [Hedin *et al.*, 1983]. This procedure starts early enough, corresponding to tangent altitudes well above 220 km, to obtain an amount of at least 40 spectra taken outside the atmosphere. They define the reference Sun spectrum. A regression is calculated on each pixel of these recorded spectra to remove the small intensity drift observed. This operation reduces the systematic errors, and results in a higher signal-to-noise ratio. For tangent altitudes lower than 220 km, the atmospheric transmittances are calculated by dividing the spectrum recorded at the current time by the reference Sun spectrum.

[13] The wave number calibration, i.e., converting pixel number of the detector to wave number using the position of known solar lines, includes a correction for the Doppler effect, as the speed of the satellite projected on the line of sight may vary during an occultation, but also from one orbit to the other. Indeed, the speed of the satellite often gets close to ± 10 km/s, depending of the relative position of the Sun and VEX, which may result in slight displacements of the solar lines, used for the calibration.

[14] When SOIR points toward the Sun, the position of the slit with respect to Venus is set in order to have its spatial direction parallel to the limb of Venus, when the

tangent altitude reaches 60 km. The slit should remain as parallel as possible to the limb in order to have the smallest vertical resolution for each separate measurement. Since the VEX satellite is set into an inertial mode while performing a solar occultation, the slit rotates slightly, and does not remain parallel to the limb during the whole occultation. This tilting angle remains small, reaching maximum values in some occultations of about 10° , which slightly decreases the vertical resolution.

[15] The orbit of Venus Express is very eccentric with its apocenter, located above the north polar region, at an altitude of 180 to 250 km above the ground. The pericenter, above the south polar region, reaches 65,000 km. For this reason, the vertical resolution of the measurement, which depends on the distance of the satellite to the limb, will also be dependent on the latitude of the measurement: high northern latitude (60° to 90°) measurements have a very good vertical resolution, the layer thicknesses are small (from 200 m to 700 m); measurements at lower latitudes (-30° to 60°) have an average vertical resolution ranging from 700 m to 2 km; southern latitudes (-90° to -30°) have a very poor vertical resolution, with layer thicknesses greater than 2 km, up to 5 km.

3.1. Method Description

[16] The analysis code is based on the ASIMUT program [Vandaele *et al.*, 2008] and has been adapted in Matlab in order to improve the retrieval procedure, i.e., by allowing the determination of the temperature, and allowing more flexibility. The algorithm uses the onion peeling approach.

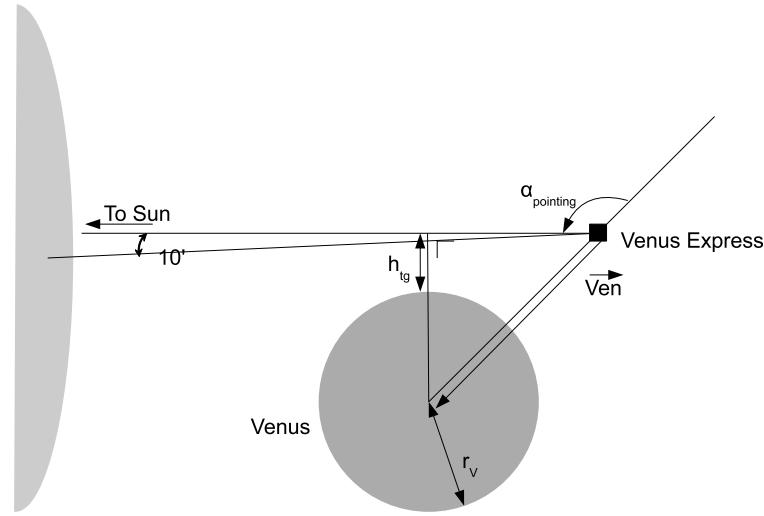


Figure 2. Pointing angles and distances definition. The distance between Venus' surface and VEX is R_{VV} . The tangent altitude is H_{tg} . The angle between the boresight of SOIR and the line linking the center of Venus and VEX is denoted α . The SOIR instrument is pointing 10' above the center of the Sun to ensure that the diffracted Sun remains within the slit for longer during the occultation [Mahieux et al., 2008].

This simplifies the problem, as for a given layer the physical properties only depend on the layers located below itself [Russel and Drayson, 1972]. The method will be described for the retrieval of the vertical density profile of a given molecule. In the cases of carbon dioxide and carbon monoxide retrievals, we will show how information can be obtained about molecular rotational temperature. The atti-

tude of VEX is calculated using the Spice for Matlab routines developed by *NASA Ancillary Information Facility (NAIF)* [2009], and used by ESA for the Venus Express mission.

3.2. Tangent Altitude and Ray Tracing

[17] The tangent altitude of each measured transmittance of one occultation series defines the atmospheric layers used in the onion peeling method, see Figure 1. The tangent

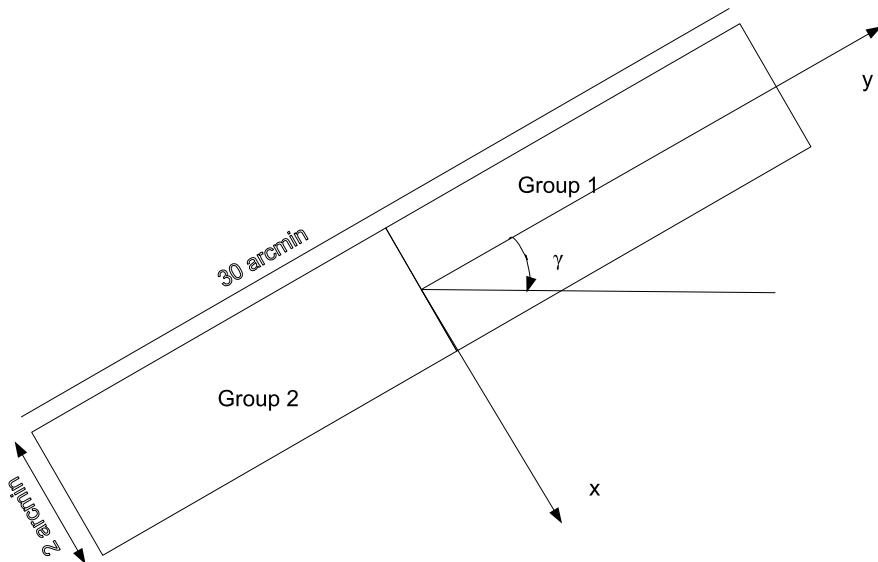


Figure 3. Representation of the slit. The slit has a width of 30' and a height of 2'. In the binning configuration 2×12 pixel rows, it is divided in two groups. The axes depicted are the ones used by the Spice routine. The slit is tilted by an angle γ . The Sun is located along the $+z$ axis. The center of the slit is the origin of the x , y , and z coordinates.

altitude is computed from the geometry of the pointing of the instrument relative to the Sun and the planet Venus as illustrated in Figure 2. The center of the projected slit is taken as a reference, see Figure 3. In this frame, the relation used to calculate the tangent altitude (H_{tg}), when no refraction is taken into account, is

$$H_{tg} = \cos(\pi - \alpha) \cdot R_{VV} \quad (1)$$

where R_{VV} is the distance from the satellite to the center of Venus and α is the pointing angle to the center of the projected slit, taking into account the 10' depointing to the center of Sun, see Figure 2. The ray tracing calculation allows the determination of the tangent height and of the absorption path in each of the layers, taking into account the refraction of the light beam into the atmosphere. The refraction model is based on the measurements performed during the Magellan mission. The refraction is calculated as described by Vandaele et al. [2008]. The determination of the radiation path through the atmosphere, i.e., the path followed by the radiation reaching the instrument, requires that the planet's curvature and refraction be taken into account. The model is based on the ray-tracing program FSCATM [Gallery et al., 1983]. Some modifications were included in the implementation, essentially dealing with the determination of the index of refraction of the air of the planet. Integrated total density, integrated partial density of the molecules fitted, mean local temperature, integrated local pressure and path length are computed for each layer. They are also corrected by accounting for the height of the projected slit in the atmosphere at each measurement.

[18] As already stated, due to telemetry limitations, the detector is divided into several groups of lines, also called bins. The pointing directions are calculated for each of these bins separately. In our example, case 2×12 , there are 2 bins, each made up of 12 detector rows. The size of the whole slit being 30', along the spatial direction, by 2', along the spectral direction, the center of bin 1 is slightly depointed depending on the tilting angle γ , see Figure 3.

3.3. Absorption Cross Section and Transmittance Calculation

[19] Molecular absorption cross sections (ACS) in each layer are calculated by a line by line procedure, extensively described by Vandaele et al. [2008], using spectroscopic parameters from Hitran 2008 [Rothman et al., 2009] adapted for the Venusian CO₂-rich atmosphere. The absorption cross sections depend on the partial and total pressures, as well as on the temperature in each layer. The transmittance due to molecule i in the layer j is calculated using

$$T_{i,j} = \exp[-n_{i,j} \cdot s_j \cdot ACS(n_{i,j}, t_j, p_j)] \quad (2)$$

where $n_{i,j}$ is the density of the molecule i in the layer j , s_j is the path length within the layer j , and t_j and p_j are the temperature and the pressure in the layer j , obtained as explained in section 3.2 from the ray tracing algorithm. The transmittance $T_{L,Molecular}$ due to the molecular species present along the whole path corresponding to the tangent height L is finally obtained by multiplying the absorbance of all molecules in all the above layers, the most outside layer

being layer 1:

$$T_{L,Molecular} = \prod_{j=1}^L \prod_{i=1}^{n_{molecules}} T_{i,j}. \quad (3)$$

[20] Aerosols, which are present in the Venus atmosphere in large quantities [Esposito, 1983; Lane and Opstbaum, 1983], also have a characteristic signature. In the infrared region probed by SOIR their signature is large, but broad and without fine structures [Wilquet et al., 2009]. Their influence on the observed spectra is to decrease the mean transmittance level. This effect is more pronounced as the line of sight goes deeper into the atmosphere.

[21] The impact of aerosols is considered by representing their contribution to the transmission by a second degree polynomial on the wave number. For layer L , the aerosol contribution can be written as

$$T_{L,Aerosols} = a_L + (\omega_n - \omega_{n_0}) \cdot b_L + (\omega_n - \omega_{n_0})^2 \cdot c_L \quad (4)$$

where a_L , b_L and c_L are the polynomial coefficients, ω_n is the wave number, ω_{n_0} is the central wave number of the scanned order. The aerosol extinction can be calculated from the parameters a_L [Bingen et al., 2003]. The results have already been discussed by Wilquet et al. [2009], and will be deeper analyzed in a future publication (V. Wilquet et al., Aerosol characterization and climatology observed by the SOIR instrument on board Venus Express, manuscript in preparation, 2010).

[22] Finally, the synthetic spectrum I_L , which will be compared to the one recorded by the instrument, is the solar spectrum attenuated by the effect of the atmospheric molecular species and aerosols, convolved by the instrumental function

$$I_L = conv \left(\underbrace{I_{Sun} \cdot T_{L,Aerosols} \cdot T_{L,Molecular}}_{\text{synthetic spectrum}}, \underbrace{\text{Gauss}(ResolSOIR)}_{\text{instrument function}} \right) \quad (5)$$

with I_{Sun} the Sun absorption cross section and $\text{Gauss}(Resol_{SOIR})$ the instrument function with the spectral resolution function of SOIR [Mahieux et al., 2008].

[23] Even if the resolution of SOIR is good enough to resolve most of the atmospheric absorption lines, it remains a few factors larger than the width of the pressure broadened atmospheric absorption lines. Therefore the pressure cannot be deduced from the measured spectrum. This is mathematically illustrated by the convolution product in equation (5): the synthetic spectrum has a FWHM of about 0.001 cm⁻¹ to 0.01 cm⁻¹, while the resolution of SOIR is about 0.15 cm⁻¹.

3.4. Order Addition

[24] The bandwidth of SOIR's AOTF, used for the diffraction order selection, is larger than the free spectral range of the echelle grating. To correctly simulate the spectra measured by the spectrometer, the contribution of the adjacent orders also has to be taken into account. Usually, the two first adjacent orders and the central order are considered to form most of the observed transmittance. In some cases, however, when very strong lines are present in

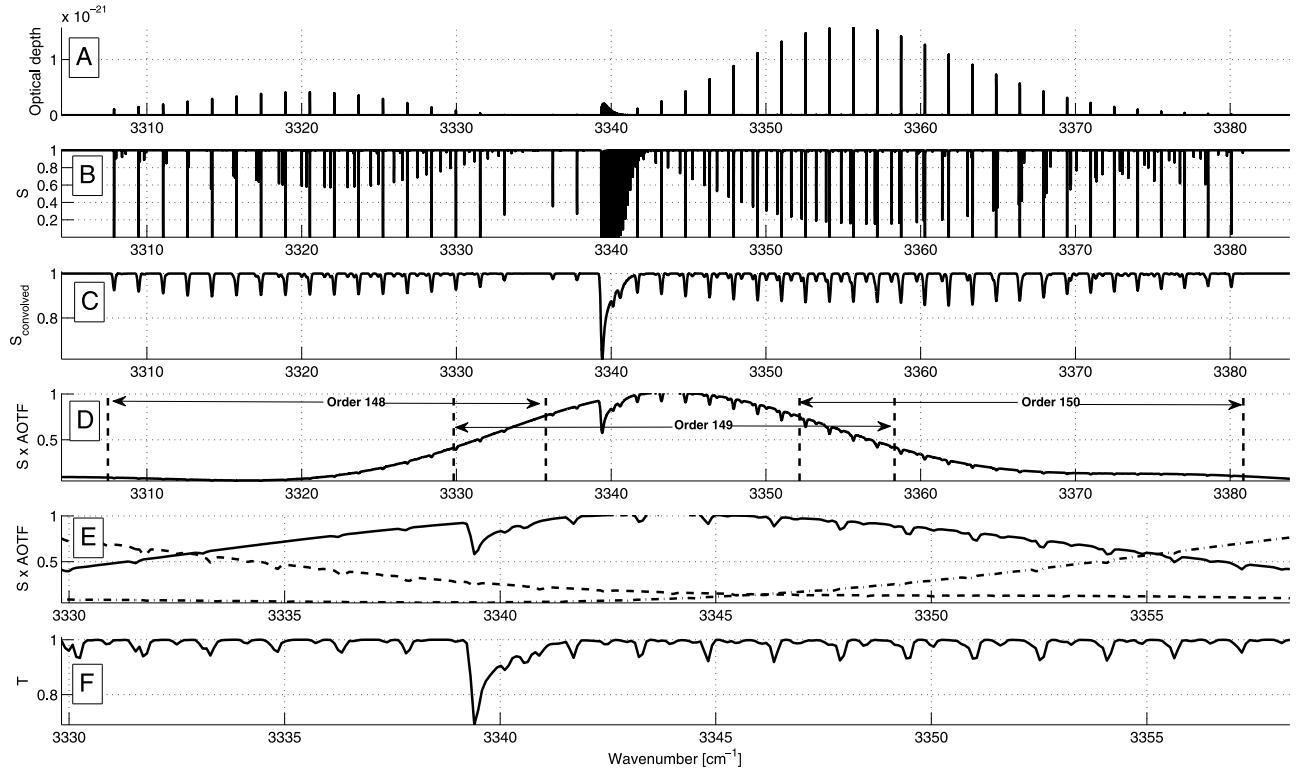


Figure 4. Procedure to simulate a SOIR spectrum. (a) The absorption cross section of CO₂ around 3340 cm⁻¹. It has been calculated for one layer with atmospheric conditions at 90 km. (b) The transmittance T of these lines after using the Beer-Lambert law. (c) The transmittance T after convolution by the instrument function. (d) The spectrum multiplied by the AOTF transfer function. The different diffraction orders are also indicated. The order overlap is clearly seen. (e) The contribution of the different orders: the bold line is the central order 149, the dashed line is the order +1, and the dashed-dotted line is the order -1. (f) The sum of the contribution of the three considered orders. One can see that the role of the adjacent orders cannot be neglected, as absorption features of the adjacent order may have significant contributions.

adjacent orders located further away from the central order, it may be necessary to consider a larger number of adjacent orders. Usually, 3 diffraction orders are taken into account. The contribution from the adjacent orders is calculated using

$$T = \frac{\sum_{i=-n_{orders}}^{n_{orders}} AOTF_i \cdot I_L}{\sum_{i=-n_{orders}}^{n_{orders}} AOTF_i \cdot I_{Sun}} \quad (6)$$

where T is the transmittance, n_{orders} is the number of adjacent orders considered, $AOTF_i$ is the values of the AOTF transfer function in the order i and I_i is the convoluted absorbance of the molecules and aerosols absorbing in order

i [Mahieux et al., 2009]. The whole procedure for constructing a synthetic spectrum is depicted in Figure 4.

4. Retrieval Algorithm

[25] The optimal estimation (OE) method developed by Rodgers [Rodgers, 1990, 2000] has been implemented to inverse the observed transmittances. All transmittances, corresponding to one occultation series considered after the selection procedure, are inversed in one go.

4.1. Implementation of the OE Algorithm

[26] The transmittance in a given layer L is given by where T_L is the transmittance at the layer L , $AOTF_p$ is the

$$T_L = \frac{\sum_{p=-n_{orders}}^{n_{orders}} AOTF_p \cdot \text{conv} \left\{ I_{Sun} \cdot T_{L,Aerosols} \cdot \prod_{j=1}^L \prod_{i=1}^{n_{molecules}} \exp[-n_{ij} \cdot s_j \cdot ACS_{ij}(n_{ij}, t_j, \omega_n j)], \text{Gauss}(\sigma_{SOIR}) \right\}}{\sum_{p=-n_{orders}}^{n_{orders}} AOTF_p \text{conv}\{I_{Sun}, \text{Gauss}(\sigma_{SOIR})\}} \quad (7)$$

$$T_{L,Aerosols} = a_L + (\omega_n - \omega_{n_0}) \cdot b_L + (\omega_n - \omega_{n_0})^2 \cdot c_L$$

$$\omega_n' j = \omega_n + d_j$$

modified AOTF transfer function, adapted on the n_{AOTF} points and translated of a value d_{AOTF} , d_i is the wave number shift of layer j and σ_{SOIR} is the resolution of the instrument. In this equation, n_{ij} , t_j , a_L , b_L , c_L , $AOTF_p$, d_{AOTF} and d_j are unknown variables.

[27] In order to fit T_L to the measured spectra, the problem is written in the form

$$y = F(x, b) + \varepsilon \quad (8)$$

where y is the measurement vector (the measured transmittances), x is the state vector (all unknown variables to be retrieved, f.e. the vertical density profiles), b represents the additional parameters used by the forward model, whose function is F . The forward function F describes the complete physics of the measurement, including the description of the instrument. F is the set of L functions T_L described in equation (7). ε represents the noise on the measurement.

[28] Let us consider the retrieval of the vertical profiles of $n_{molecules}$ from a series of $n_{spectra}$ transmittances acquired during one occultation. The $n_{spectra}$ spectra have defined a structured atmosphere composed of $m_{layer} = n_{spectra}$ layers (see Figure 1), whose boundaries correspond to the tangent heights of each measurement. The state vector contains all the variables fitted during the retrieval procedure. These variables are the a , b and c parameters of the aerosol contribution ($3 \times m_{layer}$ variables), and the logarithm of the molecular densities of the targeted species in all the defined layers ($n_{molecules} \times m_{layer}$ variables). The temperature in each layer may also be fitted (m_{layer} variables). Even if the line intensities of all molecules are temperature dependent, the temperature is only fitted for molecules having regularly spaced absorption lines or absorption lines close to each others, such as CO or CO₂. Moreover a wave number shift can be introduced and fitted for each spectrum (m_{layer} variables) to improve the matching of the absorption lines positions from the spectroscopic database and the observed position. The AOTF transfer function may be fitted too: First, a slight variation is allowed on several points, usually around 40, to improve the quality of the fit (n_{AOTF} variables); Second, the AOTF transfer function may also be translated along the wave number axis, to correct for possible deviations of the wave number – AOTF frequency relation [Mahieux et al., 2008] (1 variable). It leads to a total number of variables of $n_{var} = (5 + n_{molecules}) \cdot m_{layer} + n_{AOTF} + 1$.

[29] The Rodgers formalism solves the problem written in equation (8) by starting from an a priori value of the variable x iterating on the following equation:

$$\begin{aligned} x_{k+1} &= x_k + (S_a^{-1} + K_k^T S_\varepsilon^{-1} K_k)^{-1} \\ &\cdot [K_k^T S_\varepsilon^{-1} (y - F(x_k)) - S_a^{-1} (x_k - x_a)] \end{aligned} \quad (9)$$

where x_k are the state vector values at iteration k , S_a is the covariance matrix of the fitted parameters, S_ε is the measurement error matrix, K_k is the Jacobian of the problem at iteration k , y contains the measured spectra, F is the forward model of equation (8) and x_a is the a priori state vector.

[30] The covariance matrix S_a contains the allowed variability of each variable:

$$S_a = \varepsilon \left\{ (x - x_a)(x - x_a)^T \right\} \quad (10)$$

where x_a is the a priori state vector and x is the solution of the problem. Non diagonal terms are added for the covariance of the densities, temperature, wave number shift and AOTF transfer function variations. A Gaussian dependence is used to account for correlations between the different points where the AOTF transfer function is fitted and to ensure a final smooth AOTF transfer function. They are defined as

$$S_a^\omega(i,j) = \sqrt{S_a^\omega(i,i) \cdot S_a^\omega(j,j)} \cdot \exp \left[-\left(\frac{\omega_{n_i} - \omega_{n_j}}{l_\omega} \right)^2 \right] \quad (11)$$

where S_a^ω is the submatrix of S_a corresponding to the AOTF transfer function fit, ω_{n_i} is the wave number position of point i , l_ω is the correlation length [Vandaele et al., 2006]. The correlation length has been chosen to be a fifth of the diffraction order width for the AOTF transfer function. In the case of the densities, temperature and wave number shift, the coefficients are

$$S_a^{n.t,d}(i,j) = \sqrt{S_a^{n.t,d}(i,i) \cdot S_a^{n.t,d}(j,j)} \cdot \exp \left[-\left(\frac{z_i - z_j}{l_z} \right)^2 \right] \quad (12)$$

where $S_a^{n.t,d}$ is the submatrix of S_a corresponding to the density, temperature or wave number shift variables, z_i is the altitude of point i and l_z is the correlation length, defined as the height of the projected slit at the tangent point. In the case of the aerosols contribution only diagonal terms are considered. S_a is a square matrix of size n_{var} .

[31] The measurement error matrix S_ε contains the estimated errors calculated from the signal-to-noise ratio. S_ε is not square and has a size of $n_{\omega_n} \cdot m_{layer}$, where n_{ω_n} is the number of wave number points in each interval used for the fit.

[32] The Jacobian matrix K contains the first derivatives of the forward model with relation to all the state vector variables. The derivatives

$$\frac{\partial T}{\partial a_L}, \frac{\partial T}{\partial b_L}, \frac{\partial T}{\partial c_L}, \frac{\partial T}{\partial n_{ij}}, \frac{\partial T}{\partial t_j}, \frac{\partial T}{\partial d_j}, \frac{\partial T}{\partial n_{AOTF}}, \frac{\partial T}{\partial d_{AOTF}} \quad (13)$$

are easily computed. They are all derived analytically, except for the displacement derivatives $\frac{\partial T}{\partial d_j}$, which are calculated numerically.

[33] The sub matrices of K , $\frac{\partial T}{\partial n_{ij}}$, $\frac{\partial T}{\partial t_j}$ and $\frac{\partial T}{\partial d_j}$, are inferior triangular matrices because of the onion peeling procedure.

$\frac{\partial T}{\partial d_{AOTF}}$ is a vector, as $\frac{\partial T}{\partial a_L}$, $\frac{\partial T}{\partial b_L}$, $\frac{\partial T}{\partial c_L}$ which have nonzero values only for a given layer. $\frac{\partial T}{\partial n_{AOTF}}$ is a full matrix. K has a size of $n_{var} \times (n_{\omega_n} \cdot m_{layer})$. The vectors x_k , x_a , $F(x_i)$ and y contain the state vector variables at step k , the a priori values of the variables x_k , the calculated transmittance at step k , and the $n_{spectra}$ observed spectra, respectively. The vectors x_k , x_a have a length of n_{var} , while the vectors $F(x_i)$ and y have a length of $(n_{\omega_n} \cdot m_{layer})$.

[34] The convergence of the algorithm is achieved when the following criterions are both satisfied

$$\begin{aligned} (x^t - x^{t-1}) \cdot (S_a^{-1} + K' \cdot S_\varepsilon^{-1} \cdot K)^{-1} \cdot (x^t - x^{t-1}) &< 0.001 \cdot n_{var} \\ (F^t - F^{t-1}) \cdot (S_\varepsilon \cdot K \cdot S_a \cdot K' \cdot S_\varepsilon) \cdot (F^t - F^{t-1}) &< 0.001 \cdot n_{\omega_n} \cdot m_{layer} \end{aligned} \quad (14)$$

where x^t and F^t are the state vector and the function described in equation (10), respectively.

[35] The error calculation is done using the relations

$$\begin{aligned}\hat{x} - x &= (A - I_n) \cdot (x - x_a) + G \cdot \varepsilon \\ G &= [(S_a^{-1} + K^T \cdot S_\varepsilon^{-1} \cdot K) \cdot K^T \cdot S_\varepsilon^{-1}]^{-1} \\ A &= G \cdot K\end{aligned}\quad (15)$$

where $\hat{x} - x$ is the error vector, x is the solution at convergence and \hat{x} is the exact solution, I_n is a unity matrix having the same size as A and ε is the error vector used to build S_ε . The covariance of the error vector is built by considering two sources, namely the a priori covariance S_s and the retrieval noise covariance S_m :

$$\begin{aligned}S_s &= (S_a^{-1} + K^T \cdot S_\varepsilon^{-1} \cdot K)^{-1} \cdot S_a^{-1} \cdot (S_a^{-1} + K^T \cdot S_\varepsilon^{-1} \cdot K)^{-1} \\ S_m &= (S_a^{-1} + K^T \cdot S_\varepsilon^{-1} \cdot K)^{-1} \cdot K^T \cdot S_\varepsilon^{-1} \cdot K \cdot (S_a^{-1} + K^T \cdot S_\varepsilon^{-1} \cdot K)^{-1}\end{aligned}\quad (16)$$

4.2. Fitting Procedure

[36] For a given occultation and a given diffraction order, the spectroscopic parameters of the different species to retrieve must be supplied, as well as the wave number range on which they will be fitted. More than one isotopologue for each species may be fitted simultaneously. The altitude range is automatically determined, but may also be imposed by the user.

[37] The algorithm returns the fitted values of all the parameters defined earlier. The user has the choice to use the a priori atmospheric parameters derived from the atmosphere model, or to use the ones derived from another fitting procedure, for example a quasi simultaneous CO₂ inversion.

4.3. Predetermination of the Useful Altitude Range

[38] We have developed an automatic procedure to determine the useful altitude range of the retrieval of one given occultation. Usually, atmospheric spectra are recorded at tangent altitudes starting around 220 km and down to 60 km. However, the spectra corresponding to the higher altitudes usually do not contain any absorption features, but provide a good indication of the noise level of the measurements, and spectra corresponding to the lowest altitudes are often saturated due to too high atmospheric absorption as detailed in the following.

[39] The first spectrum considered for the retrieval corresponds to the emergence of the monitored absorption structures. Only the wave number range on which the retrieval is performed is considered. The maximum (I_{\max}), minimum (I_{\min}) and mean (I_{mean}) values of the transmittance are computed on this window for each atmospheric spectrum. The depth of the strongest observed line (I_{diff}) is computed by subtracting the minimum value from the maximum value ($I_{\text{diff}} = I_{\max} - I_{\min}$). At high altitudes, usually above 180 km, just before the emergence of the absorption features, only noise is measured, and I_{diff} provides a good estimate value for it. A signal to noise value ($\overline{SNR}_{180-220}$) representative of these high altitudes is cal-

culated by considering the average values of the signal and of the noise between 220 km and 180 km, $\overline{I_{\text{mean}}}$ and $\overline{I_{\text{diff}}}$,

$$SNR_{180-220} = \frac{\overline{I_{\text{mean}}}}{\overline{I_{\text{diff}}}}. \quad (17)$$

[40] The first spectrum to consider for the retrieval procedure corresponds to the altitude when the absorption structures are seen above the noise. From trials, a factor of 0.002 has been introduced, thus the first spectrum is defined for

$$\frac{I_{\text{mean}}}{I_{\text{diff}}} < 0.002 \cdot SNR_{180-220} \quad (18)$$

From this altitude and below, spectral structures are seen in the spectra.

[41] The lower bound of the altitude range is defined when saturation occurs: the atmospheric absorption lines have transmittances reaching zero values at their center, before convolution by the instrument resolution function. From this altitude, the information contained in the spectra becomes useless, and reliable retrievals cannot be made.

[42] The criteria used for the selection of first and last spectra are clearly demonstrated when plotting the quantity $\frac{I_{\text{mean}}}{I_{\text{diff}}}$ as a function of the altitude, as illustrated in Figure 5. The slope radically changes at two well defined altitudes. Values of $\frac{I_{\text{mean}}}{I_{\text{diff}}}$ are almost constant above the altitude of point 2 on Figure 5c corresponding to the appearance of the absorption features as illustrated by Figures 5a and 5b, corresponding to one spectrum without absorption structures and one with them, respectively. The altitude of the last spectrum to be included in the retrieval, point 3 on Figure 5c, is obtained at the second change in slope. Below this altitude, most of the absorption lines are saturated – at infinite resolution – and the surface under the structures seen in the convolved spectrum recorded by the instrument are no longer proportional to the density of the absorbing species.

[43] Saturation may be observed if the transmittance of a spectral line reaches values close to zero. In that case, all the information contained in the observed spectral line is lost, as the shape of the line does not vary with increasing densities. It is not directly observed on the spectra measured by SOIR, as the resolution of the instrument is much larger than the width of the observed spectral lines. Detection procedures of such events have been developed, to avoid meaningless density profiles. All absorption lines having a transmittance (at infinite resolution) lower than 0.15 are considered as saturated (J. van der Auwera, private communication, 2010). The transmittance before convolution and order addition $\tilde{T}_L(\omega_n)$ is defined as

$$\tilde{T}_L(\omega_n) = \prod_{j=1}^L \prod_{i=1}^{n_{\text{molecules}}} \exp[-n_{ij} \cdot s_j \cdot ACS_{ij}(n_{ij}, t_j, \omega_{n_i})]. \quad (19)$$

[44] If saturation occurs, two actions are taken: (1) all the points ω_n^* for which we have $\tilde{T}_L(\omega_n^*) < 0.15$ are removed from the fitted wave number range and (2) the spectrum of layer j^* at iteration k of the Rodgers algorithm is said to be

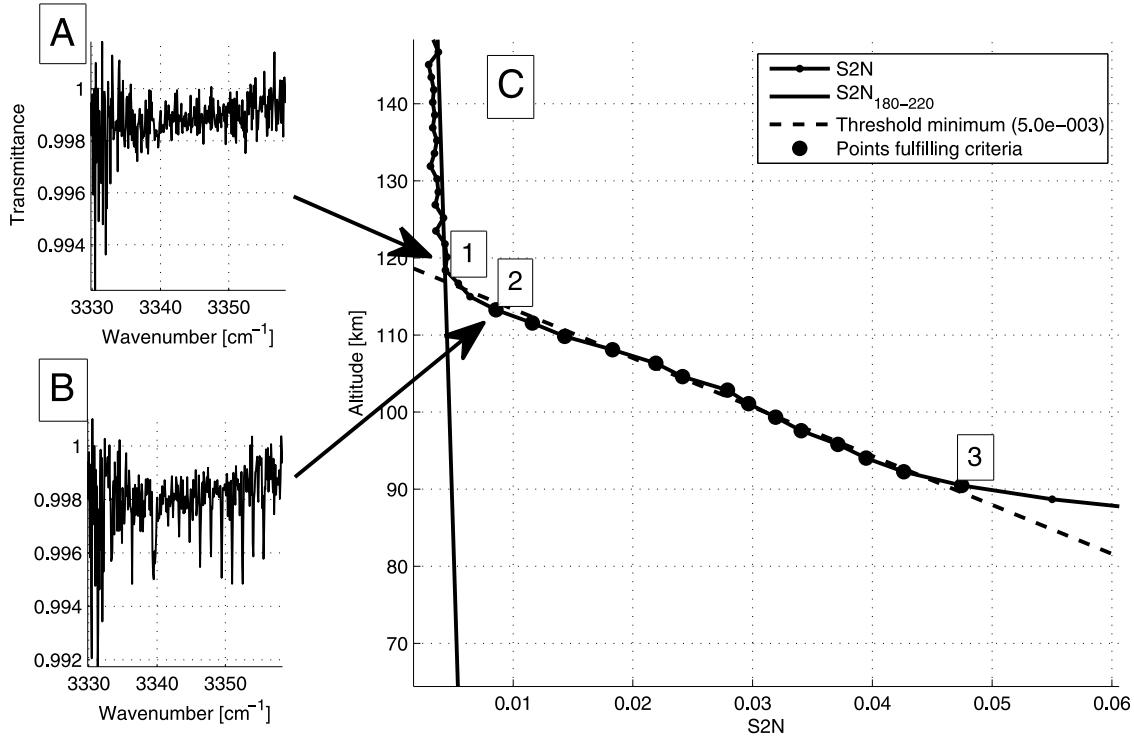


Figure 5. Determination of the altitude range. The altitude range is determined by studying the quotient $I_{\text{mean}}/I_{\text{diff}}$, with I_{mean} being mean value of the spectrum on the interval fitted and I_{diff} being the maximum less the minimum value on the interval. Figure 5c depicts its variation as a function of altitude. At high altitude, above point 1 (see spectrum in Figure 5a), $I_{\text{mean}}/I_{\text{diff}}$ is a majoring value of the noise level. When the absorption structures start appearing the spectra, at point 2 (see spectrum in Figure 5b), an inflection point is observed in the curve. It is observed that the second inflection point (point 3) corresponds to the altitude when the absorption lines are completely saturated.

saturated if more than 40 percent of the absorption lines coming from the central order are saturated, or 40 percent of the detector is covered by saturated lines. In that case, the altitude range is modified to suit the saturation criterion: all the layers j located under and including the current layer j^* are excluded from the retrieval, i.e., for $j^* \leq j \leq m_{\text{layer}}$. To that end, the corresponding lines and columns are removed from the S_a , K , S_e , Y , F , x_k and x_a matrices or vectors. That condition is evaluated at each step of the OE algorithm.

[45] However, the aerosol background can be retrieved for the layers in which saturation occurs, if the saturated lines are removed from the wave number fitted range. The variables a , b and c of the aerosol background are thus retrieved on the entire occultation, while the density, temperature and displacement are only retrieved in the non saturated altitude range.

4.4. Model of the Atmosphere

[46] In this work, we have used temperature and pressure vertical profiles from the VIRA model for altitudes up to 100 km [Seiff *et al.*, 1985]. For higher altitudes (from 140 km and upward), data were taken from the model of Hedin *et al.* [1983] as suggested by Mueller-Wodarg and Tingle [2008]. The transition between the two data sets was performed by splines interpolating the temperature and reconstructing the pressure through the hydrostatic law.

4.5. Temperature Dependence

[47] SOIR spectra are sensitive to the atmospheric temperature through its influence on the strength of the absorption features of, for example, CO or CO₂. Temperature will influence the overall shape of one specific vibrational band, as shown in Figure 6. Each peak belongs to a given rotational transition J . If temperature increases the position of the band absorption maximum will move from low J lines to higher J lines, as the distribution of line intensities for different rotational transition changes. From the overall shape, temperature can then be derived even if some of the information contained in this shape disappears because of the curvature of the AOTF transfer function, as illustrated in Figure 6.

[48] This approach leads to the determination of the rotational temperature of the gas in one vibrational band. Local thermodynamical equilibrium (LTE) needs to be satisfied in order to assimilate the rotational temperature to the kinetic temperature, which may not be necessarily the fact [Gilli *et al.*, 2009]. However, the non-LTE effects only weakly affect ground state rotational transition bands (M. A. Lopez-Valverde, private communication, 2010).

[49] The shape of the AOTF function will thus influence the temperature retrieval, as a good knowledge of the function is needed in order to correctly build the synthetic spectrum. However, the correction of the AOTF function using the n_{AOTF} points as described in section 4.1 will only

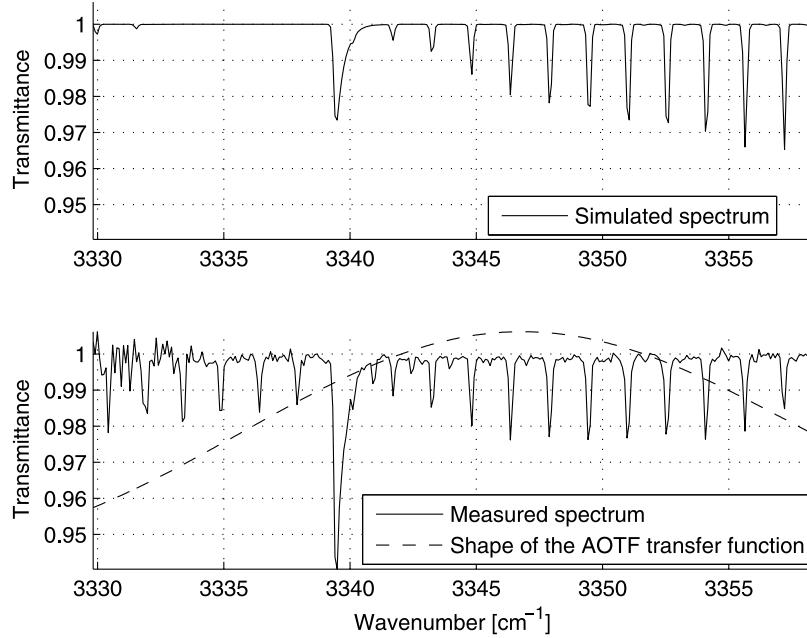


Figure 6. Temperature dependency. The temperature information is influenced by the shape of the AOTF transfer function. (top) A CO₂ spectrum in diffraction order 149, simulated for a layer (atmospheric conditions: altitude of 110 km, pressure of 0.152 Pa, temperature of 181.2 K, total density of 0.1×10^{-4} mol/m³, volume mixing ratio of 0.965, path length of 200 km) but without order addition nor contribution of the AOTF transfer function. It is clearly seen that the maximum of absorption of the band is located around 3347 cm⁻¹. (bottom) The measured spectrum and the shape of the AOTF function. It is obvious that the maximum of the measured transmittance has been displaced closer to the maximum of the AOTF transfer function.

locally modify the transfer function, and thus not the overall shape.

4.6. Wave Number Scale Definition

[50] Using the theoretical relations of the echelle grating, a linear relationship between the wave number, the order and the pixel number can be defined [Schroeder and Hilliard, 1980; Nevejans *et al.*, 2006; Mahieux *et al.*, 2008]

$$\begin{cases} d(\sin \theta_i + \sin \theta_m) = m \cdot \lambda \\ \lambda = 10^{-2}/\omega_n \\ pixel = f(\theta_m) \end{cases} \Rightarrow pixel = f \left[\arcsin \left(\frac{m \cdot 10^{-2}}{d \cdot \omega_n} - \sin \theta_i \right) \right] \quad (20)$$

where θ_i is the incident angle on the grating, θ_m is the refracted angle on the grating, m is the diffraction order, λ is the wavelength, d is the groove spacing of the echelle grating and f is a linear function linking the diffracted angle and the pixel number.

[51] The calibration has been performed by comparing the positions of selected solar lines observed by SOIR outside the Venus atmosphere with a reference atlas of solar lines [Altieri *et al.*, 2009], taking into account a correction applied to the observed solar lines for the Doppler effect induced by the motion of Venus Express relative to the Sun. Indeed, the satellite is orbiting around Venus at a speed which is varying significantly. The speed at apogee may reach absolute values of 10 km/s relative to Venus. The speed vector has to be projected onto the line of sight of the SOIR instrument.

The sign of the speed may be positive or negative, depending of the satellite movement direction relative to the Sun.

$$\omega'_n = \omega_n \cdot \left(1 + \frac{v_{VEX}}{c} \right) \quad (21)$$

where ω_n is the wave number position of the line – taken from the solar atlas, v_{VEX} is the speed of the Venus Express satellite relative to Venus projected onto the line of sight of the instrument, c is the speed of light and ω'_n is the wave number corrected for the Doppler effect. The rotation of the planet itself is not taken into account, as its speed is very low: about 1.811 m/s at the equator. The speed of the planet Venus around the Sun projected onto the line of sight may be neglected during an occultation, as the eccentricity of the orbit of Venus is very small (~0.00677323).

[52] After these corrections, it has been observed that small variations remain when considering the coincidence of the observed and database absorption lines. The justification comes from small temperatures variations within the instrument that may induce such variations to the above described calibration. The residual is fitted using the relation

$$\omega'_n = \omega_n + (e + f \cdot (\omega_n - \omega_{n0})) \quad (22)$$

where e and f are the coefficients of a first-order polynomial. These two variables are fitted on the whole occultation,

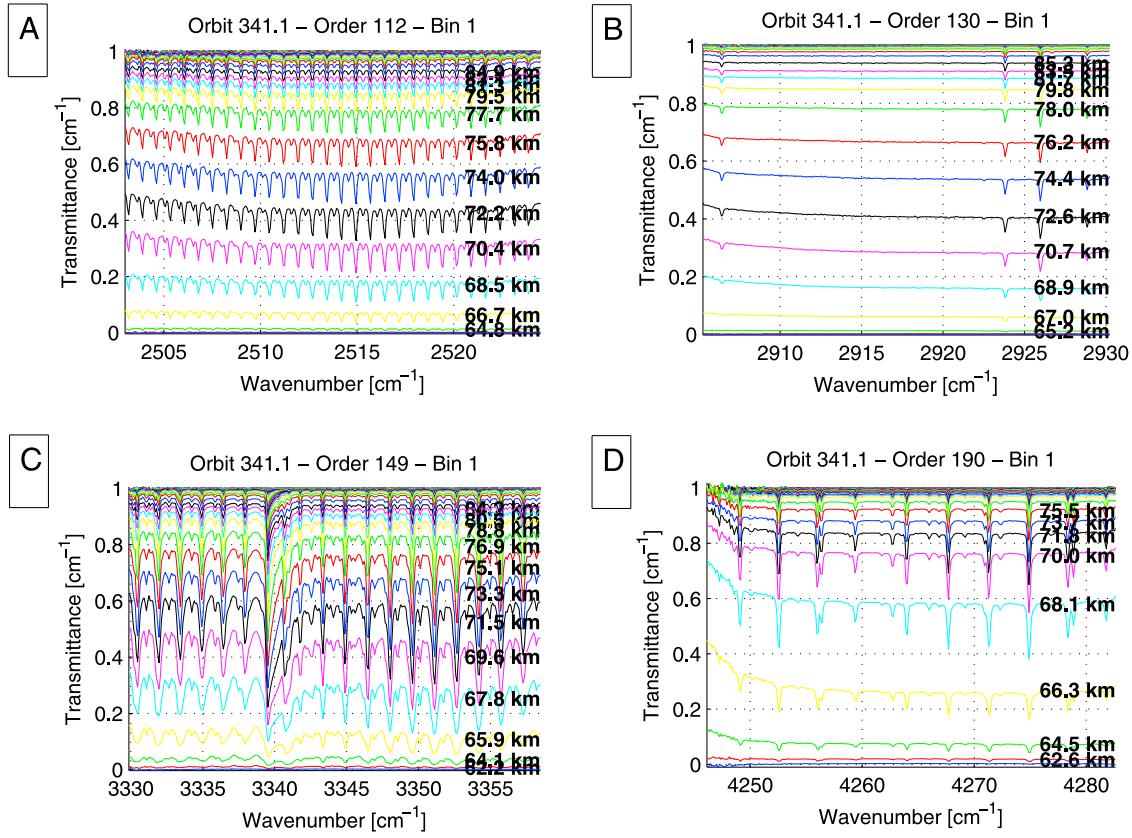


Figure 7. Spectra measured during orbit 341.1 (28 March 2007). The selected orders were (a) 112, (b) 130, (c) 149, and (d) 190. Each image depicts all the spectra taken during the occultation in the selected diffraction order. The absorption features observed in Figure 7a come from $^{16}\text{O}^{12}\text{C}^{18}\text{O}$, in Figure 7b from $^1\text{H}^{15}\text{Cl}$ and $^1\text{H}^{15}\text{Cl}$, in Figure 7c from $^{16}\text{O}^{12}\text{C}^{16}\text{O}$, and in Figure 7d from $^{12}\text{C}^{16}\text{O}$.

meaning they have the same value for all m_{layer} spectra. For this reason, they are independent of the wave number shifts of each layer d_i .

4.7. Measurement Error and Pointing Error

[53] Two external independent sources of errors, coming from different sources are considered in the present work: the measurement error and the satellite pointing error. The first one, the measurement error calculation, has been described by Mahieux *et al.* [2008]. The signal-to-noise ratio obtained is very good, reaching average values of around 500 up to 3000.

[54] The second one, the satellite pointing error, is obtained directly from ESOC (<ftp://ssols01.esac.esa.int/pub/data/GDP/VEX/ATTDEV/>), the European Space Operation and Control of the European Space Agency, which provides the pointing angle error, from which the altitude error is computed. The largest error values at the limb are of the order of 500 m for orbits when the satellite is far away of the planet, and less than 50 m while the satellite is close to the planet.

5. Examples of Applications

[55] In the following we will illustrate the method developed above with two different retrieval cases. The focus is placed on the analysis of the method and its merits.

Further scientific discussion on the results, considering a more complete data set, will be pursued in the second part of this series.

5.1. Carbon Dioxide Retrieval

[56] One order per occultation has been systematically dedicated to carbon dioxide measurements. The diffraction order 149 ($3329.8\text{--}3358.3\text{ cm}^{-1}$) corresponds to a spectral interval in which the main isotopologue $^{16}\text{O}^{12}\text{C}^{16}\text{O}$ presents a relatively strong absorption band. As mentioned by Vandaele *et al.* [2008], this band allows us to cover tangent altitudes from about 75 km up to 120 km. But saturation of the lines generally restricts the minimum altitude to approximately 100 km. To simulate this order, the adjacent orders ± 1 and ± 1 are taken into account, ranging from 3307.47 cm^{-1} to 3380.86 cm^{-1} . The resolution of the instrument in this region is 0.16 cm^{-1} . The three most important bands in intensity, which are included in the simulation, are the $21102\text{--}00001$, $22202\text{--}01101$ and $30003\text{--}01101$ transitions, having maximal intensities of $9.1 \times 10^{-22}\text{ cm}^{-1}/(\text{molecule cm}^{-2})$, $4.4 \times 10^{-24}\text{ cm}^{-1}/(\text{molecule cm}^{-2})$ and $1.3 \times 10^{-24}\text{ cm}^{-1}/(\text{molecule cm}^{-2})$, respectively, comprising 306 lines in total. Figure 7 shows the coverage of this order measurement.

[57] A complete description of the retrieval of orbit 341 (28/03/2007) is given in the following. Table 2 summarizes the orbit information, while Figure 8 depicts all the spectra measured in the four scanned diffraction orders.

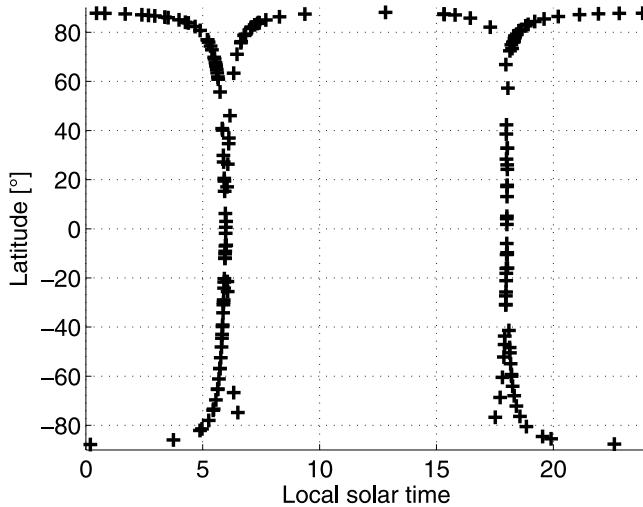


Figure 8. Trace of the measurement points of diffraction order 149 on a local solar time versus latitude map. The measurements occur either at 0600 LT or 1800 LT, on the morning or the evening terminators. A gap is seen at the morning terminator between 50° and 60° north, as well as on the evening terminator between 35° and 70° north. This is due to the orbit geometry, implying that few solar occultations occur for VEX at these latitudes.

The measurements were taken between 0620:01 LT and 0629:54 LT, during an egress case of the satellite. The latitude at the tangent altitude of 65 km was 82.05° , and the longitude was 348.20° . The local solar time was 17.30. The speed of the satellite projected on the line of sight was

-8.76 km/s. The useful altitude range for the retrieval extended from 102.7 km up to 120.0 km, resulting from the criteria developed before. The vertical profiles for the CO₂ density and the temperature are given in Figure 9. Figure 10 illustrates the quality of the fit, which is performed on only a portion of the recorded spectra (green curve). The layers below 102.7 km have been removed because some of the lines were saturated, and were not fulfilling the criteria described in section 4.3. The averaging kernels for the main fitted parameters are presented in Figure 11. As expected, the averaging kernels corresponding to the background parameters (a , b , and c , not shown for the latter) are only defined at the level to which they correspond.

[58] The averaging kernels or the CO₂ concentrations have maximum values increasing while going deeper in the atmosphere, because there is more spectroscopic information present in the spectra: the path length is longer and the density is higher for low-altitude measurements. It may also be observed that the highest-altitude kernels are also defined at the surrounding altitudes, corresponding to the fact that they influence as well the layers located below them. The averaging kernel for the wave number shift is much more intricate, as a shift in a layer will partially influence all the other layers. No direct conclusion can be drawn in this case.

[59] The error covariances matrices have been calculated for the baseline parameters a , b and c , the CO₂ density, the temperature and the AOTF fitted values. The error covariances of a , b , c and wave number shift parameters are diagonal matrices, as the parameter of a given layer does not depend on the other layers parameters. The influence of the other layers on the density parameters is observed in its error covariance matrix by showing nonzero values

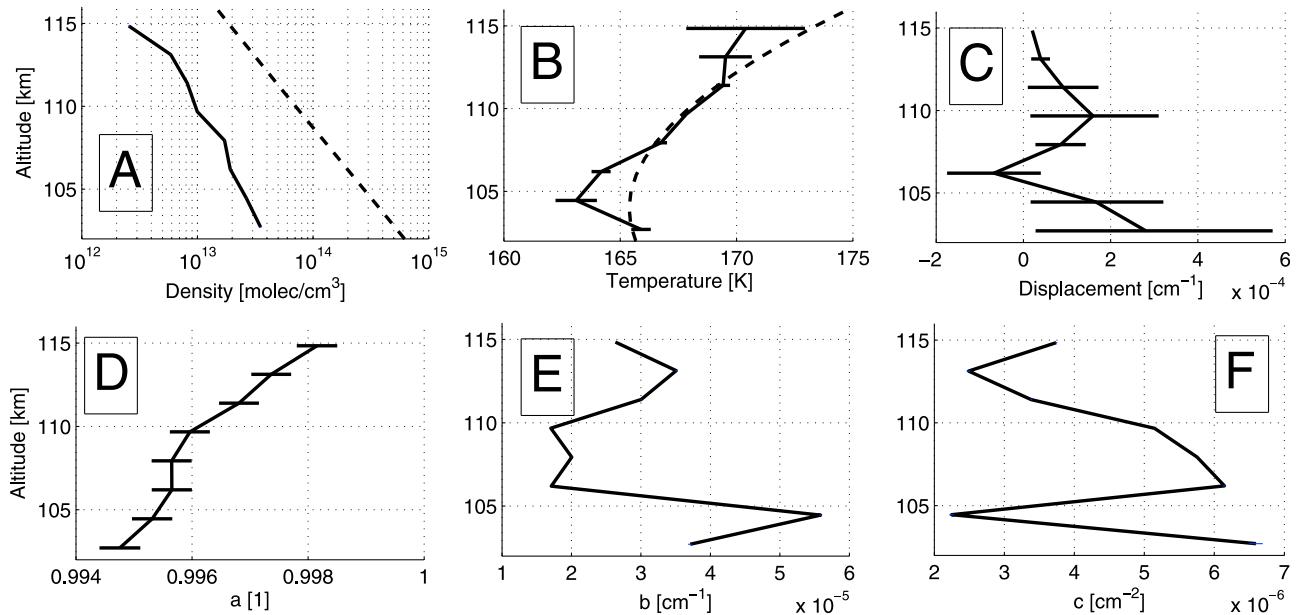


Figure 9. Vertical profiles obtained from the retrieval of the orbit 341, using order 149: (a) CO₂ density, (b) temperature, (c) wave number shift, and (d–f) aerosols parameters. Reference vertical profiles (dashed), from the combination of the VIRA model [Seiff et al., 1985] and the VTS3 model [Hedin et al., 1983]; retrieved profile (plain). The errors obtained from the optimal estimation algorithm are also plotted on the graphs.

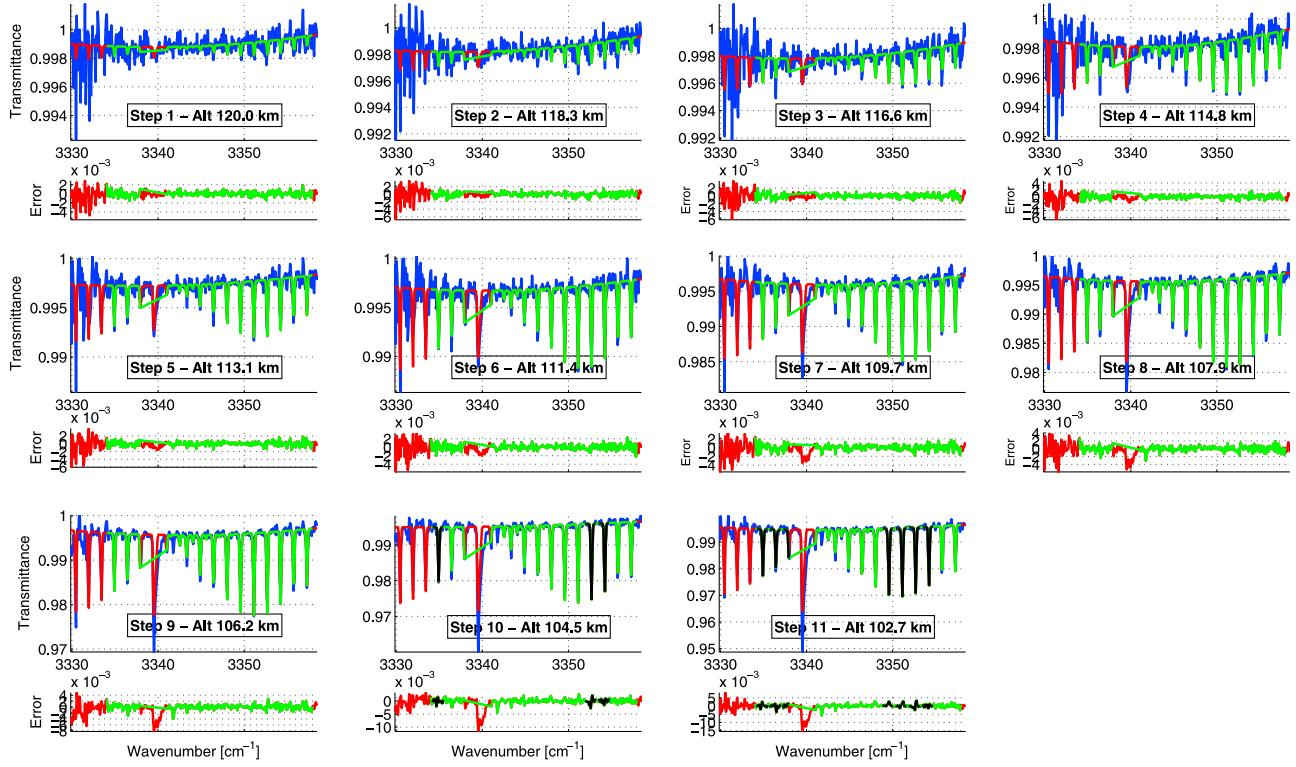


Figure 10. Illustration of the quality of the spectral fit obtained during the CO_2 retrieval of orbit 341, order 149. For each altitude, (top) comparison of the observed (blue) and fitted (green) transmittances (red represents the calculated transmittance outside the fitting wave number range), and (bottom) illustration of the residuals (observed–fitted transmittances).

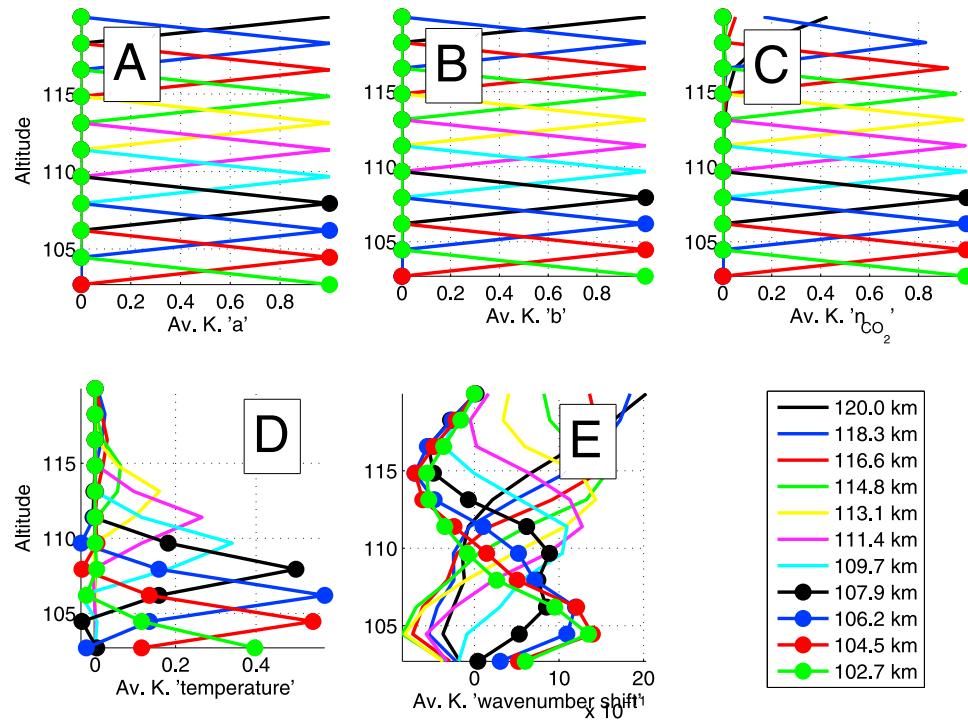


Figure 11. Averaging kernels of the last step of the optimal estimation algorithm from the retrieval of CO_2 of orbit 341, order 149: (a) aerosols a parameter, (b) aerosols b parameter, (c) CO_2 vertical densities, (d) rotational temperature vertical profile, and (e) wave number shift. The averaging kernels show the area of dependency of each variable.

Table 2. Details of the Orbit 341.1 on 28 March 2007^a

UTC Time	Tangent Altitude (km)	Latitude (deg)	Longitude (deg)	Pointing Angle (deg)	Distance to Surface (km)	Height of the Projected Slit (km)
2007-03-28T06:25:45.880	120.0	82.7	349.2	117.9	934.0	2.0
2007-03-28T06:25:44.880	118.3	82.7	349.2	118.0	936.6	2.0
2007-03-28T06:25:43.880	116.6	82.7	349.2	118.1	939.2	2.0
2007-03-28T06:25:42.880	114.8	82.7	349.1	118.1	941.8	2.0
2007-03-28T06:25:41.880	113.1	82.6	349.1	118.2	944.4	2.0
2007-03-28T06:25:40.880	111.4	82.6	349.0	118.3	947.0	2.0
2007-03-28T06:25:39.880	109.7	82.6	349.0	118.4	949.6	2.0
2007-03-28T06:25:38.880	107.9	82.6	349.0	118.4	952.2	2.0
2007-03-28T06:25:37.880	106.2	82.6	348.9	118.5	954.8	2.1
2007-03-28T06:25:36.880	104.5	82.5	348.9	118.6	957.5	2.1
2007-03-28T06:25:35.880	102.7	82.5	348.9	118.6	960.1	2.1

^aSummary of the local UTC time, the tangent altitude, the latitude and longitude of the tangent point, the pointing angle as defined in Figure 2, the distance between the surface of Venus and Venus Express and the height of the projected slit at the tangent point.

outside the diagonal: the upper layers have an influence on the lower layers, coming from the onion peeling method. It is not observed for the error covariances of the temperature, as the temperature dependence is weak.

[60] The vertical profiles described here above cover an altitude range between 120 and 102.7 km. As explained before, the retrieval can take place in a given order when the absorption features appear in the spectra, and is stopped when saturation of the absorption lines occurs. Using other diffraction orders, other CO₂ bands with different line intensities can be investigated, allowing a wider vertical coverage. This is illustrated in Figures 12 and 13 where the densities of CO₂ and the rotational temperature obtained

using different orders in quasi consecutive orbits are presented.

5.2. Carbon Monoxide Retrieval

[61] The main isotopologue ¹²C¹⁶O of carbon monoxide presents some absorption features in diffraction order 190 (4246.1–4282.4 cm⁻¹). Two transitions are considered: the (2–0) band, with 28 lines and a maximum intensity of 4.5×10^{-21} cm⁻¹/(molecule cm⁻²) and the (3–1) band with 21 lines and a maximum intensity of 3.2×10^{-28} cm⁻¹/(molecule cm⁻²). To simulate this order, the adjacent orders +1 and -1 are taken into account, ranging from 4223.7 cm⁻¹ to 4305.0 cm⁻¹. The resolution of the instrument in this region is 0.23 cm⁻¹.

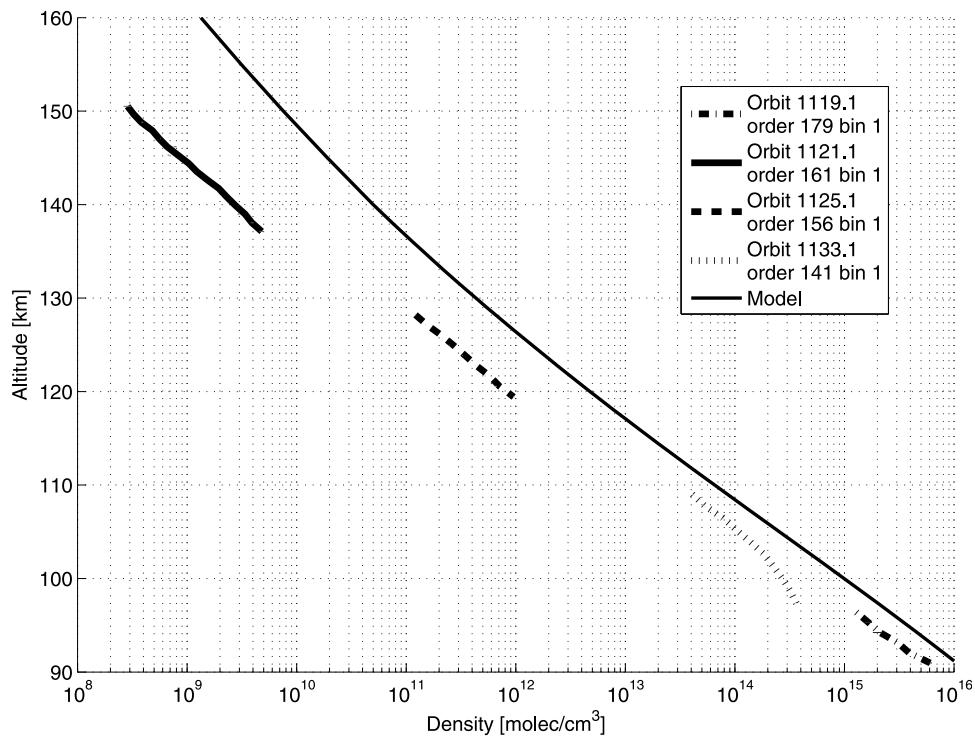


Figure 12. CO₂ vertical profiles coming from four different diffraction orders, allowing covering a wider vertical altitude range. They were obtained from the orbits 1119.1 (14 May 2009, 80.4°N, 1830 LT), 1221.1 (16 May 2009, 83.1°N, 1648 LT), 1125.1 (20 May 2009, 84.6°N, 1912 LT), and 1133.1 (28 May 2009, 87.6°N, 2200 LT).

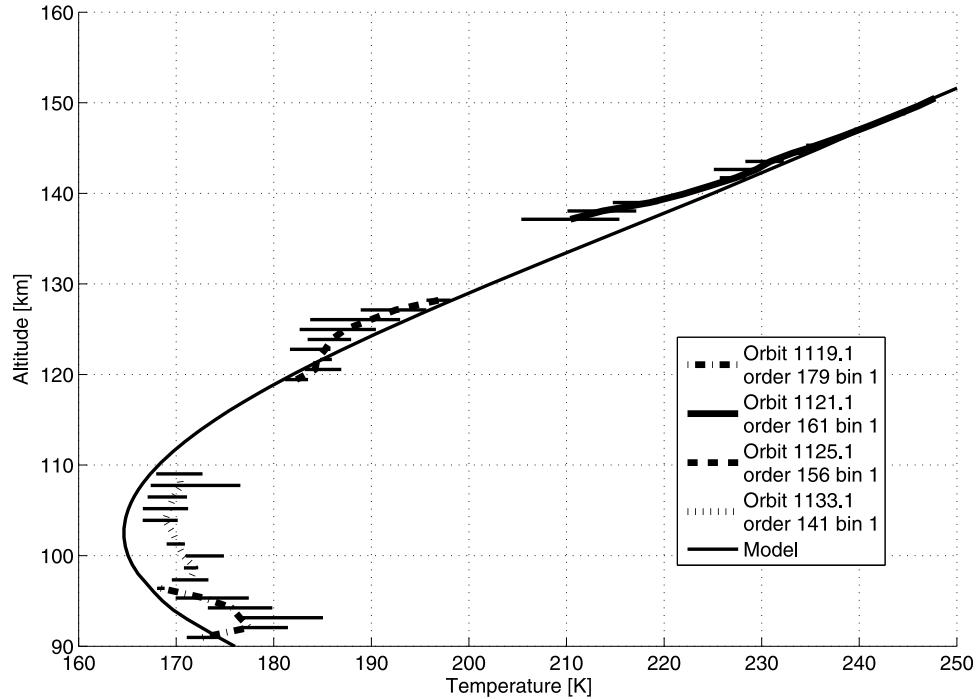


Figure 13. CO_2 rotational temperature vertical profiles coming from four different diffraction orders, allowing covering a wider vertical altitude range. They were obtained from the orbits 1119.1 (14 May 2009, 80.4°N , 1830 LT), 1221.1 (16 May 2009, 83.1°N , 1648 LT), 1125.1 (20 May 2009, 84.6°N , 1912 LT), and 1133.1 (28 May 2009, 87.6°N , 2200 LT).

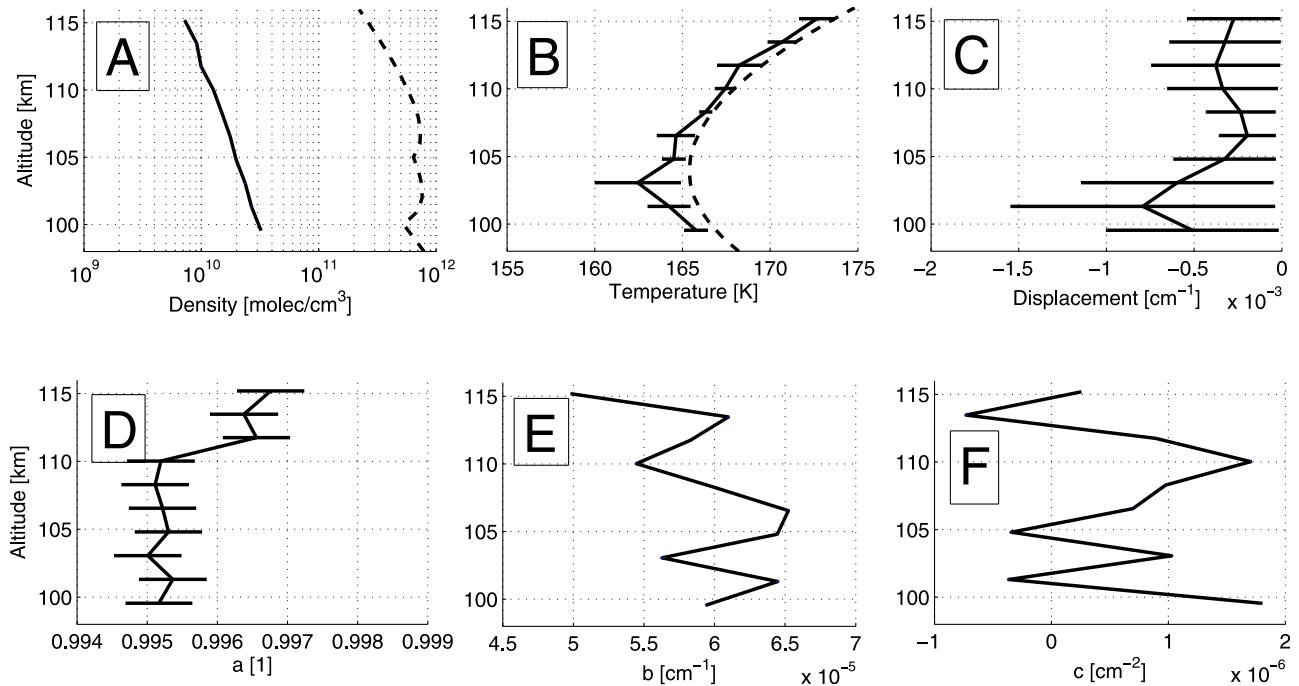


Figure 14. Vertical profiles of (a) CO density, (b) temperature, (c) wave number shift, and (d–f) aerosols parameters obtained from the retrieval of the orbit 341, using order 190: retrieved profile (plain); reference vertical profiles (dashed), combination of the VIRA model [Seiff et al., 1985] and the VTS33 model [Hedin et al., 1983]. The temperature is the one retrieved during the CO_2 retrieval. The errors obtained from the optimal estimation algorithm are also plotted on the graphs.

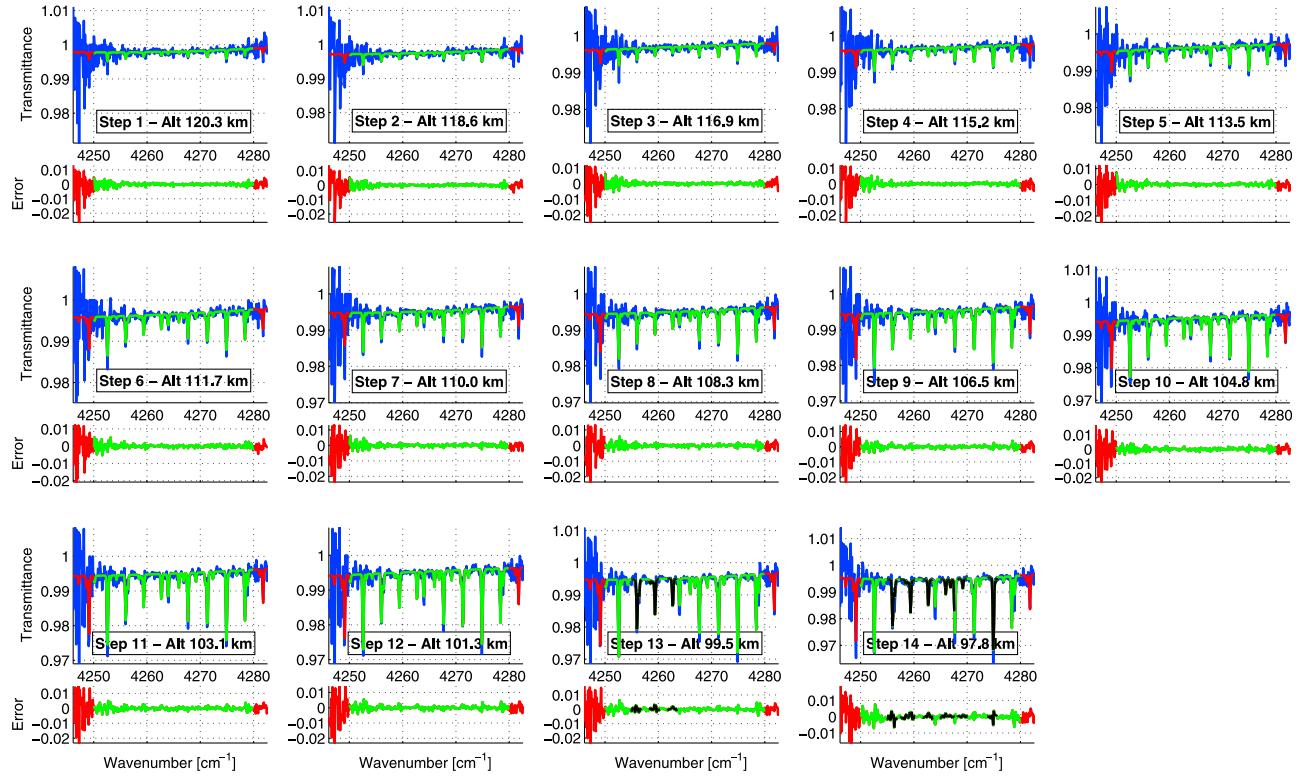


Figure 15. Illustration of the quality of the spectral fit obtained during the CO retrieval of orbit 341, order 190. For each altitude, (top) comparison of the observed (blue) and fitted (green) transmittances (red represents the calculated transmittance outside the fitting wave number range), and (bottom) illustration of the residuals (observed–fitted transmittances).

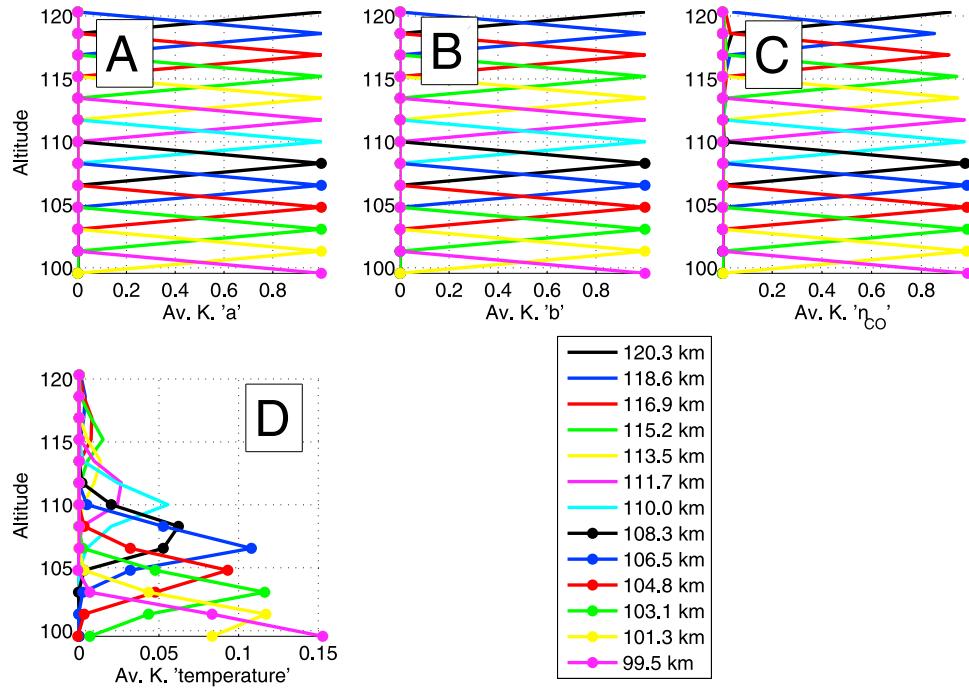


Figure 16. Averaging kernels of the last step of the optimal estimation algorithm from the retrieval of CO of orbit 341, order 190: (a) aerosols a parameter, (b) aerosols b parameter, (c) CO vertical densities, and (d) rotational temperature profile. The averaging kernels show the area of dependency of each variable.

[62] Order 190 was scanned during occultation 341, the same one used to illustrate the retrieval of CO₂. The preliminary altitude range was defined from 119 km down to 65 km, but during the fit it has been reduced to 119 km down to 99.7 km, because of the saturation occurring at the lower altitudes. Similar to what has been shown for the CO₂ results; Figures 14–16 show the results of the fitting procedure in the case of CO. As for CO₂ the concentrations in one layer is almost independent of the surrounding layers, and the information content also increase when going deeper in the atmosphere. In this example, temperature has also been retrieved and is favorably compared to the temperature retrieved from CO₂ (see Figures 9b and 14b).

6. Conclusion

[63] The present paper shows how the densities of some key constituents of the atmosphere of Venus, as well as rotational temperature profiles can be retrieved from the SOIR data, using a retrieval technique based on the OE method. This method allows the simultaneous derivation of the densities of all the target species, including different isotopologues. The total pressure and total density profiles can also be obtained using a specific procedure, in the case of carbon dioxide retrievals.

[64] The error bars on the retrieved density are lower than 6%, and on the temperature are lower than 1.5%. The method provides reliable vertical profiles that will be used to construct a database of the upper Venus atmosphere at the terminator, and will allow us to study the temporal and latitudinal variations of the composition of the atmosphere. A CO₂ climatology is being built and will be described and discussed in the forthcoming companion paper. At the same time, a detailed inventory of the composition of the atmosphere is being done in terms of CO, H₂O/HDO, HF, HCl, or SO₂.

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